Riemannian geometric aspects of Penrose-type inequalities

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Variational Formulations of Mass Inequalities PMT and Riemannian Penrose Inequality Generalizations Answers to the Questions

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Generalizations Answers to the Questions

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Known Mass Inequalities

The Positive Mass Theorem (Schoen-Yau 1979, Witten 1981)

Among all time-symmetric asymptotically flat initial data sets for the Einstein-Vacuum Equations, flat Euclidean 3-space is the unique minimizer of the total mass.

 $m \ge 0$

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The Penrose Inequality (Huisken-Ilmanen 2001, Bray 2001)

Among all time-symmetric asymptotically flat initial data sets for the Einstein-Vacuum Equations with an outermost minimal surface Σ of area *A*, the Schwarzschild slice is the unique minimizer of the total mass.

$$m \geq rac{1}{2}R \quad (R = \sqrt{A/4\pi})$$

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Summary

The Angular Momentum Case

Is the Kerr slice the unique minimizer of the total mass among all asymptotically flat axisymmetric maximal gauge initial data sets for the Einstein-Vacuum Equations with an outermost minimal surface Σ of area *A* and (Komar) angular momentum *J*?

$$m \ge \frac{1}{2} \left(R^2 + \frac{4J^2}{R^2} \right)^{1/2}$$

Natural Questions

The Charged Case I

Is the Reissner-Nordström slice the unique minimizer of the total mass among all asymptotically flat time-symmetric initial data sets for the Einstein-Maxwell Equations with an outermost minimal surface Σ of area *A* and charge *Q*?

$$m \geq \frac{1}{2}\left(R + \frac{Q^2}{R}\right)$$
 $(Q = \int_{S^2_{\infty}} E \cdot n)$

Recall;

The Time-Symmetric Einstein-Maxwell Contraints

$$S_g=2(|E|_g^2+|B|_g^2), \quad \operatorname{div}_g E=\operatorname{div}_g B=0, \quad E imes_g B=0$$

Natural Questions

The Charged Case II

(Gibbons '84) Is the Majumdar-Papapetrou slice with the horizon consisting of two components of opposite charges the unique minimizer of the total mass among all asymptotically flat time-symmetric initial data sets for the Einstein-Maxwell Equations with an outermost minimal surface $\Sigma = \bigcup \Sigma_i$ of area $A = \Sigma A_i$ and charges $\{Q_i\}$?

$$m \geq \sum_{i} \frac{1}{2} \left(R_i + \frac{Q_i^2}{R_i} \right)$$

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Charged Case I

Answers:

 Yes, provided Σ is connected (Jang 1979, Huisken-Ilmanen 2001).

Recall;

The Topology of Horizon $\boldsymbol{\Sigma}$

 Σ is diffeomorphic to $\cup S^2$ each S^2 stable minimal surface in (M^3, g) .

Charged Case I

Answers:

- Yes, provided Σ is connected (Jang 1979, Huisken-Ilmanen 2001).
- No, in general (Weinstein-Y. 2004).

Recall;

The Topology of Horizon $\boldsymbol{\Sigma}$

 Σ is diffeomorphic to $\cup S^2$ each S^2 stable minimal surface in (M^3, g) .

Charged Case II

Answers: No : Brill-Lindquist satisfies $m < \frac{1}{2} \sum_{i=1}^{2} R_i$ (Dain-Weinstein-Y. 2010).

Remark: By setting quasi-local mass-like quantities;

$$m_i = \frac{1}{2} (R_i + \frac{Q_i^2}{R_i}), \text{ or } \sqrt{\frac{1}{4}R_i^2 + \frac{J_i^2}{R_i^2}}$$

the inequalities

$$m\geq \sum m_i.$$

do not hold, as they reduces to $m \ge \frac{1}{2} \sum R_i$ in vacuum.

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Flow Σ by Inverse Mean Curvature; $\frac{\partial x}{\partial t} = \frac{1}{H}n$.

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- Use the Geroch Monotonicity (sharpened by Jang)

$$rac{dm_{H}(\Sigma(t))}{dt} \geq rac{R}{32\pi}\int_{\Sigma(t)}\mathcal{S}_{g} \geq 0$$

of the Hawking Mass $m_H(\Sigma) = \frac{R}{2}(1 - \frac{1}{16\pi}\int_{\Sigma}H^2)$ and the scalar curvature is $S_g = 2(|E|^2 + |B|^2)$.

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- Jump over discontinuities with Huisken-Ilamen's weak flow.

Penrose Inequality with Charge when Σ is connected

$$m \ge \frac{1}{2}\left(R + \frac{Q^2}{R}\right)$$
 with '=' iff Reissner-Nordström.

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- The total asymptotic area of Σ is $A = 8\pi\mu^2$.
- The total charge is $Q = 2\mu$.

An Almost Counter-Example

$$m-\frac{1}{2}\left(R+\frac{Q^2}{R}\right)=\mu\left(2-\frac{3}{\sqrt{2}}\right)<0$$

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MP is not asymptotically flat.

An Almost Counter-Example

$$m-\frac{1}{2}\left(R+\frac{Q^2}{R}\right)=\mu\left(2-\frac{3}{\sqrt{2}}\right)<0$$

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An Almost Counter-Example

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- MP is not asymptotically flat.
- MP has no horizon.
- In Weinstein-Y.(2004), two copies of Majumdar-Papapetrouare truncated at their necks, and glued to rectify the features.

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Variational Formulations of Mass Inequalities PMT and Riemannian Penrose Inequality Generalizations Answers to the Questions

One-Black-Hole Argument for the Charged Case I Jang/Huisken-Ilmanen argument

A Two-Black-Hole Counterexample to the Charged Case I The Majumdar-Papapetrou Metric Relation to Cosmic Censorhip

A Two-Black-Hole Counterexample to the Charged Case II Brill-Lindquist initial data Property of Stable Minimal Surface

Schwarzschild/Reissner-Nordström spacelike slice as Solitons Evolving Metrics by Isometries

Summary

► Jang (1979):

$$m \geq \frac{1}{2} \left(R + \frac{Q^2}{R} \right)$$

► Jang (1979):

$$m-\sqrt{m^2-Q^2} \le R \le m+\sqrt{m^2-Q^2}$$

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Jang (1979):

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- Only the upper bound on *R* follows from Cosmic Censorship using Penrose's heuristic argument.
- Our counter-example violates the lower bound.

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$$g_{ij} = u^4 \delta_{ij}, \quad u = \left(1 + \frac{\mu}{2r_1} + \frac{\mu}{2r_2}\right)$$

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 r_1, r_2 are the distances to $p_1 = (0, 0, 1), p_2 = (0, 0, -1)$ in \mathbb{R}^3 .

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A Counter-Example (Dain-Weinstein-Y. 2010)

$$m-\frac{1}{2}\sum_{i}R_{i}<0$$

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Summary

Multiple Components

Stability of $\Sigma \Rightarrow$ multiple components

 If the outermost horizon is connected, then there exists x ∈ Σ ∩ {z = 0}. Then an estimate by Schoen says there exists ε > 0 such that

$$\sup_{\Sigma_1 \cap B(x,\varepsilon)} |A| \le C \int_{\Sigma_1 \cap B(x,2\varepsilon)} |A|^2 dx$$

where the RHS is $o(\mu)$ as $\mu \to 0$ and $\Sigma_1 = \Sigma \cap B(x, r)$ with r < 1. This says $|\Sigma_1| > C$, with *C* independent of μ

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Bray's Penrose inequality

$$2\mu \ge \sqrt{rac{|\Sigma|}{16\pi}}$$

says $|\Sigma| \rightarrow 0$ as $\mu \rightarrow 0$.

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Summary

Conformal Flows via Linear Elliptic Equations

Consider the one-parameter familes of metric on $\mathbf{R}^3 \setminus \{0\}$

$$g(t) = \mathcal{U}(t)^4 (dr^2 + r^2 d\omega^2)$$

where

$$\mathcal{U}(t) = e^{-t} + \frac{m}{2re^{-t}}$$

are Schwarzschild metrics, while

$$\mathcal{U}(t) = \left(e^{-t} + \frac{m+Q}{2re^{-t}}\right)^{1/2} \left(e^{-t} + \frac{m-Q}{2re^{-t}}\right)^{1/2}$$

are Reissner-Nordström metrics.

Conformal Flows via Linear Elliptic Equations

Set u(t) = U(t)/U(0), then a function v(t) defined by $u(t) = \exp \int_0^t v(\tau) d\tau$ satisfy, for all t,

$$L_{g(t)}v(t) = 0, \ \ L_{g(t)}v(t) = \frac{3}{4}|E|_{g}^{2}v(t)|$$

for Schwarzschild and Reissner-Nordström metrics respectively, where $L_g = \triangle_g - \frac{1}{8}S_g$. This is crucial to Bray's proof of RPI.

Question:

Any physical meaning to this gauge invariance?

Partial Answer:

Ohashi-Shiromizu-Yamada (2009)



Results For Solutions of the Einstein-Maxwell Constraints:



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 connected $\Rightarrow m \ge \frac{1}{2}\left(R + \frac{\alpha}{R}\right)$.

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$$\Sigma$$
 connected $\Rightarrow m \geq \frac{1}{2}\left(R + \frac{Q^2}{R}\right).$

► Σ not connected $\Rightarrow m - \frac{1}{2}\left(R + \frac{Q^2}{R}\right) < 0$ is possible (glued and perturbed Majumdar-Papapetroudata).

►
$$\Sigma$$
 not connected $\Rightarrow m - \frac{1}{2} \sum_{i} \left(R_{i} + \frac{Q_{i}^{2}}{R_{i}} \right) < 0$ is possible (Brill-Lindquist initial data.)

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●>



Open Problems:





Open Problems:

► The Charged Case





Open Problems:

► The Charged Case

•
$$R \leq m + \sqrt{m^2 - Q^2}$$



Open Problems:

- The Charged Case
 - $R \leq m + \sqrt{m^2 Q^2}$

The Rotating Case

Open Problems:

- The Charged Case
 - $R \leq m + \sqrt{m^2 Q^2}$

The Rotating Case

•
$$m \ge \frac{1}{2} \left(R^2 + \frac{4J^2}{R^2} \right)^{1/2}$$
 for Σ connected

Open Problems:

- The Charged Case
 - $R \leq m + \sqrt{m^2 Q^2}$
- The Rotating Case

•
$$m \ge \frac{1}{2} \left(R^2 + \frac{4J^2}{R^2} \right)^{1/2}$$
 for Σ connected
• $\frac{R^2}{2} \le m^2 + \sqrt{m^4 - J^2}$

S. Ohashi, T. Shiromizu, and S. Yamada. Riemannian Penrose inequality and a virtual gravitational collapse. Physical Review D 80 (2009) 047501.

- S. Dain, G. Weinstein and S. Yamada. A counterexample to a Penrose inequality conjectured by Gibbons. Classical and Quantum Gravity. 28 (2010) 085015.
- S. Yamada and G. Weinstein
 On a Penrose Inequality with Charge. Commun. Math.
 Phys. 257 (2004), 703–723.